

Plato WP127 workshop
25 June 2016

Galactic physics Gaia-Plato synergy

D. Katz (on behalf of the Gaia people)



Milky-Way science case

	Gaia	Plato
Galactic structure <ul style="list-style-type: none"> Density profiles Disc warp and flare Etc... 	α, δ, π	
Galactic kinematics & dynamics <ul style="list-style-type: none"> Velocity distribution functions Gravitational instabilities Resonances Radial migration 	$\mu_{\alpha}, \mu_{\delta}, V_R$	
Gas, dust & stellar phases <ul style="list-style-type: none"> Gas and dust distribution Gas infall and outflow Star formation, IMF, yields Chemistry: iron, alpha elements, r and s processes 	A_V, E_{B-V} $T_{\text{eff}}, \log g, [\text{Fe}/\text{H}], [\text{X}/\text{Fe}]$ multiplicity (mass) variability	$T_{\text{eff}}, \log g$ Evolutionary stage (e.g. RGB vs RC) variability
Milky-Way formation/evolution <ul style="list-style-type: none"> Age/Kinematic, Age/Metallicity, Age/abundances relation Star Formation Rate, mass assembly history 	Age (from "isochrones")	Age

Parallaxes: precisions

	B1V	G2V	M6V
V-Ic [mag]	-0.22	0.75	3.85
Bright stars	5-14 μas	5-14 μas	5-14 μas
V=15 mag	26 μas	24 μas	9 μas
V=20 mag	330 μas	290 μas	100 μas

α, δ, π

$\mu_\alpha, \mu_\delta, V_R$

A_V, E_{B-V}

Teff, logg
[Fe/H]
[X/Fe]

variability

Age

10%

5%

1

2

$10 \mu\text{as} \leftrightarrow 10\%$ on distances at 10kpc

$A_V = 0$

Parallaxes: e.g., Mapping the Milky-Way

William Herschel (1785)



α, δ, π

$\mu_\alpha, \mu_\delta, V_R$

A_V, E_{B-V}

$T_{\text{eff}}, \log g$
 $[\text{Fe}/\text{H}]$
 $[\text{X}/\text{Fe}]$

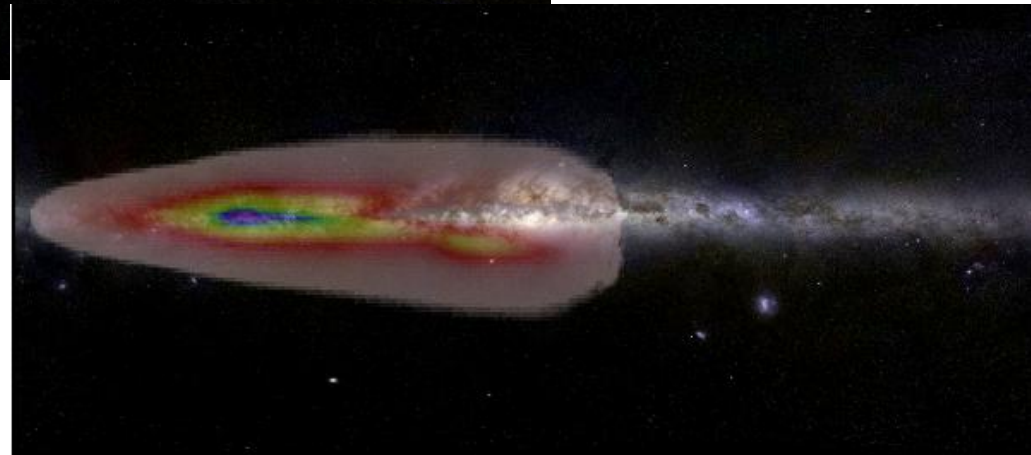
variability

Age



Hipparcos (1993)
30 000 distances
with 10% precision

Gaia
150 millions distances
with 10% precision



Radial velocity: K giant spectrum

HIP 86564 K5 $V = 6.64$

α, δ, π

μ_α, μ_δ

V_R

A_V, E_{B-V}

$T_{\text{eff}}, \log g$

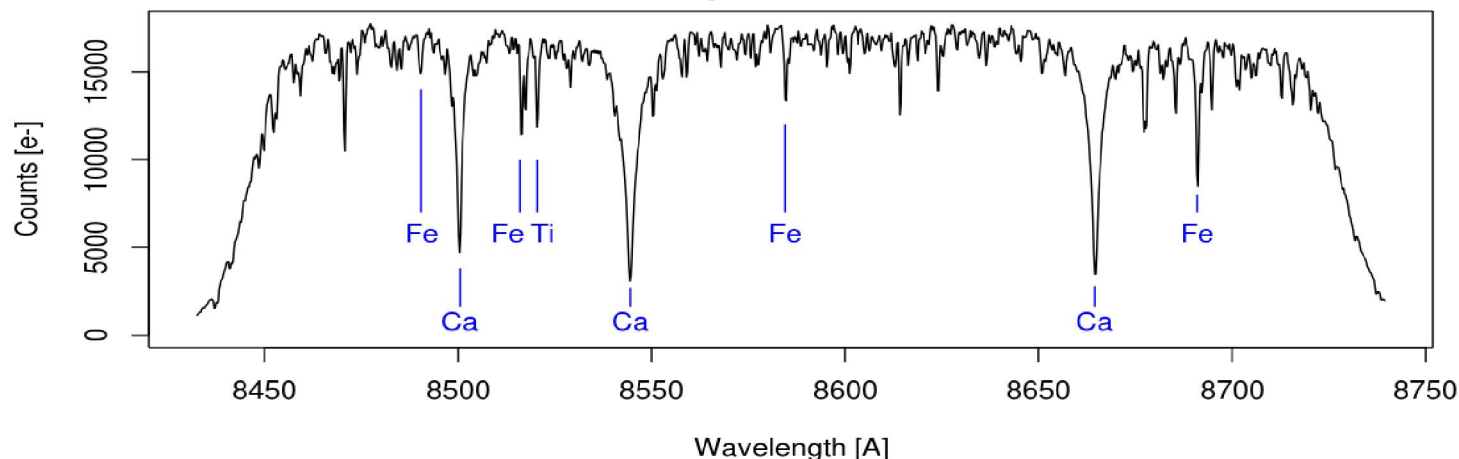
$[\text{Fe}/\text{H}]$

$[\text{X}/\text{Fe}]$

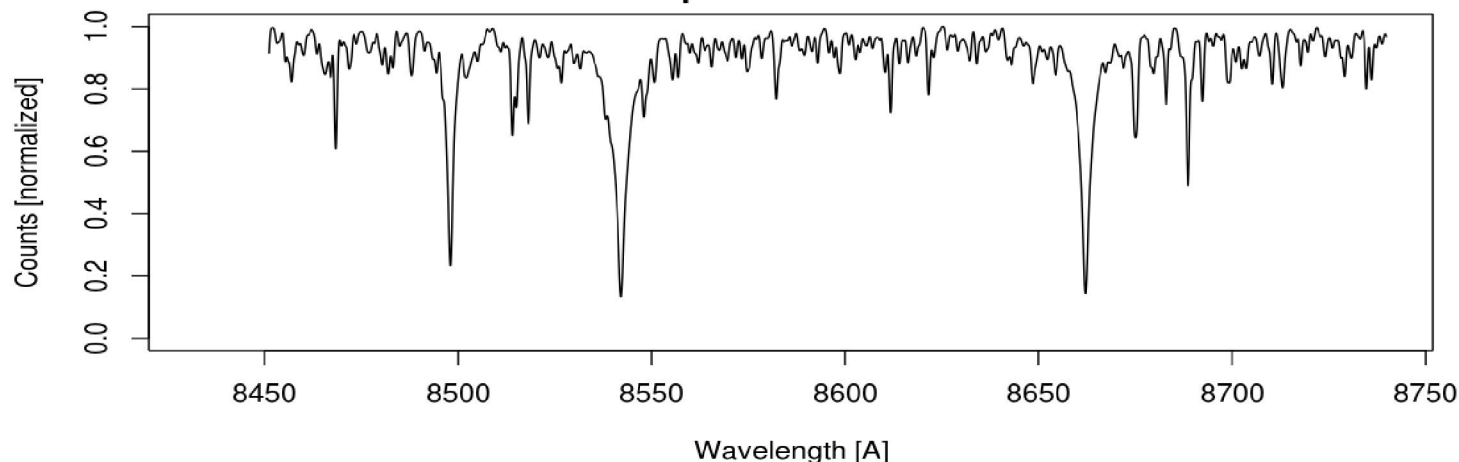
variability

Age

Gaia-RVS spectrum of HIP 86564

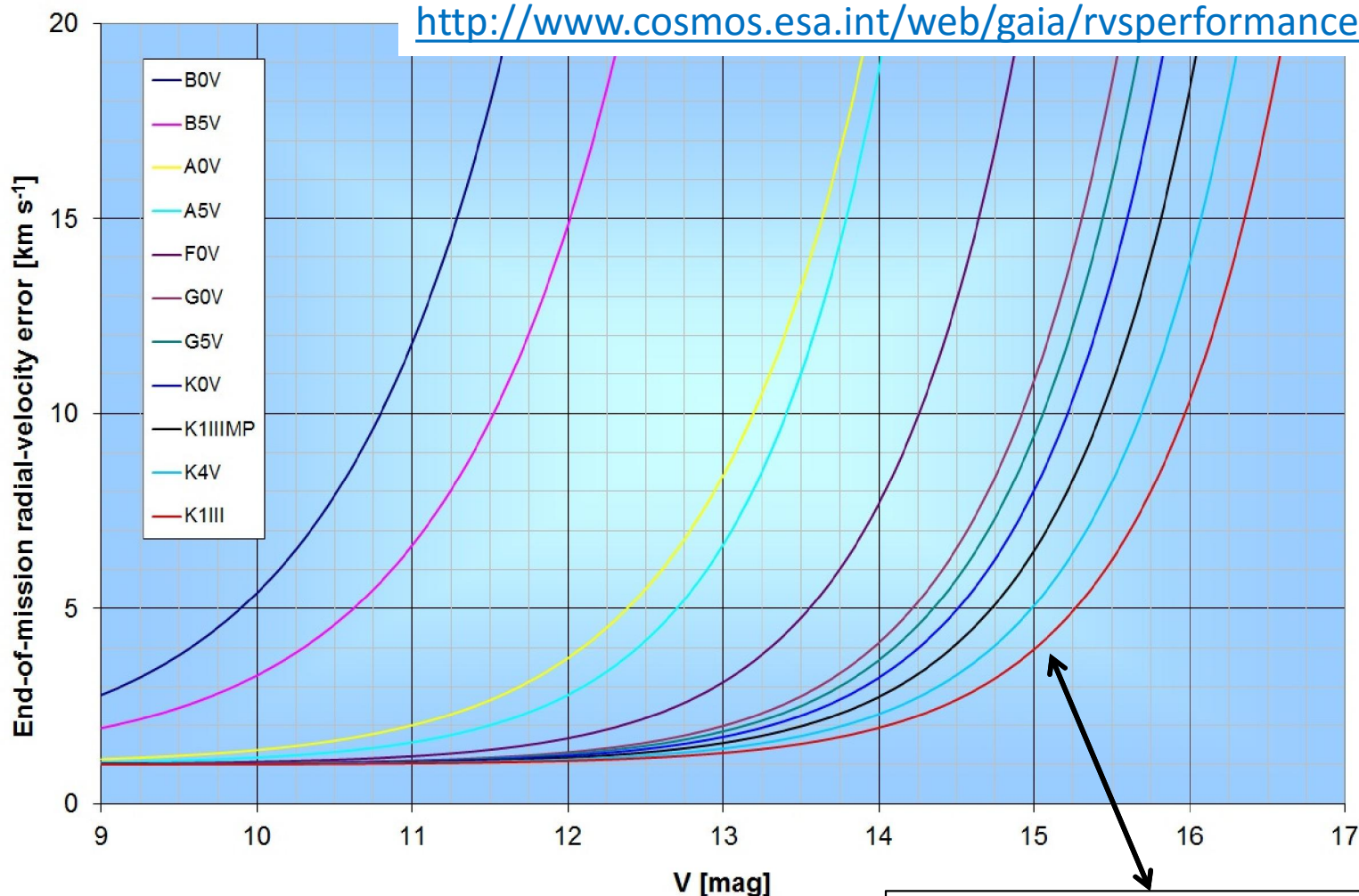


Narval spectrum of HIP 86564



Radial velocity: precisions

<http://www.cosmos.esa.int/web/gaia/rvperformance>



α, δ, π

μ_α, μ_δ

V_R

A_V, E_{B-V}

$T_{\text{eff}}, \log g$

$[\text{Fe}/\text{H}]$

$[\text{X}/\text{Fe}]$

variability

Age

K giant (solar metallicity)

- 1 kpc $\rightarrow \sigma V_R \leq 1$ km/s
- 10 kpc $\rightarrow \sigma V_R \approx 4$ km/s
- 15 kpc $\rightarrow \sigma V_R \approx 10$ km/s

A_v, E_{B-V} : BP-RP, RVS

credits: ESA/Gaia/DPAC/CU6/Observatoire de Paris-Meudon/Olivier Marchal, Carine Babusiaux & David Katz

α, δ, π

$\mu_\alpha, \mu_\delta, V_R$

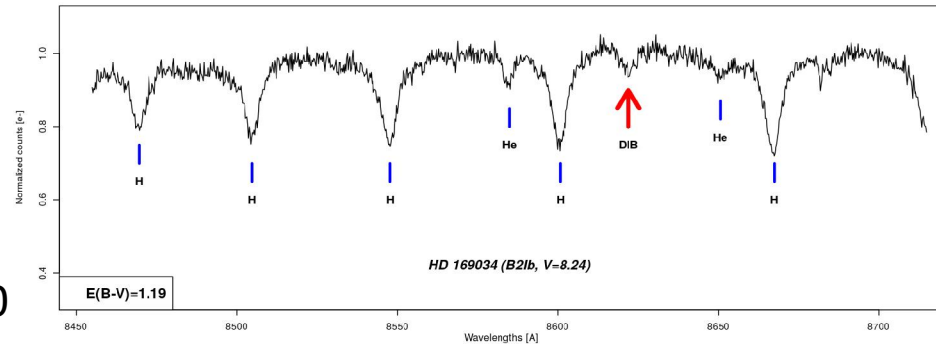
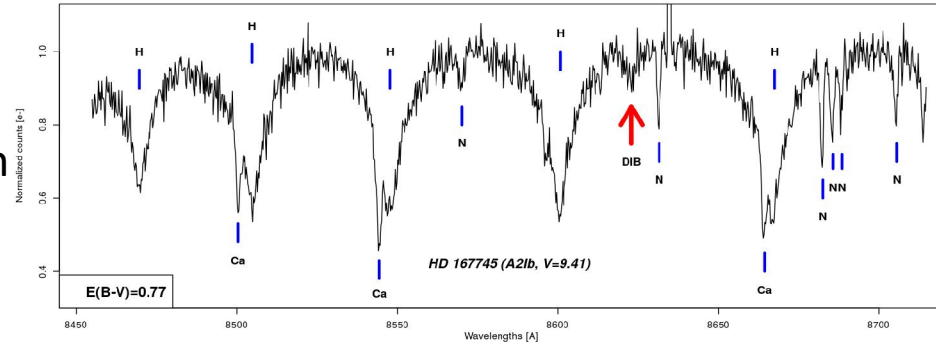
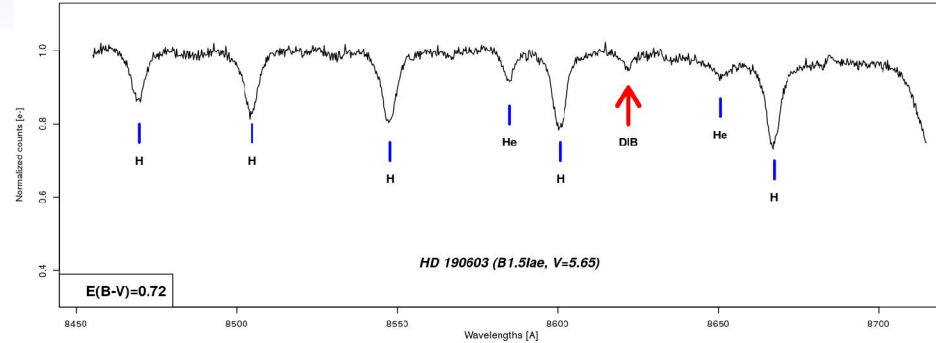
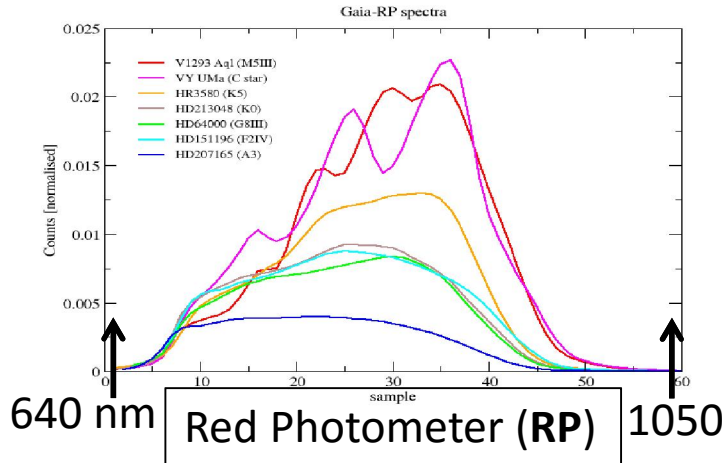
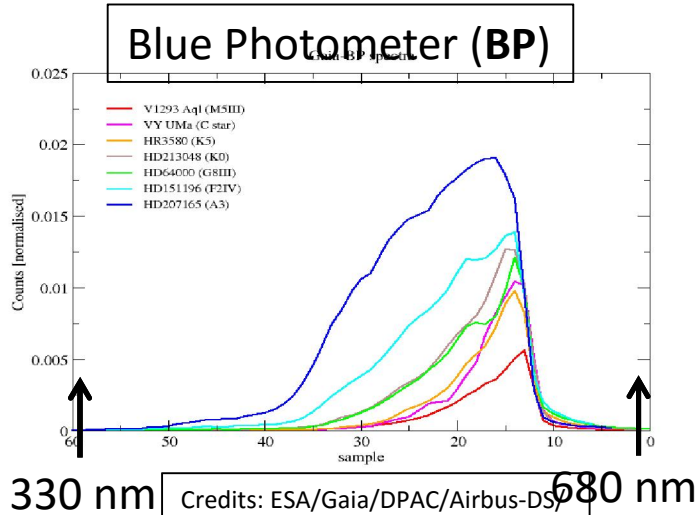
A_v ,

E_{B-V}

$T_{\text{eff}}, \log g$
[Fe/H]
[X/Fe]

variability

Age



Bailer-Jones, Andrae, Arcay et al., 2013, A&A, 559, A74

$G=15 \rightarrow \sigma A_0 = 0.1 \text{ mag}$

Diffuse Interstellar Band(DIB): 8620 Å

$E(B-V)$: $E(B-V) = 2.72 \times \text{EW (Ang)}$

Munari, Tomasella, Fiorucci et al., 2008, A&A, 488, 969

$T_{\text{eff}}, \log g, [\text{Fe}/\text{H}], [\alpha/\text{Fe}]: \text{Bp-Rp}$

α, δ, π

$\mu_{\alpha}, \mu_{\delta}, V_R$

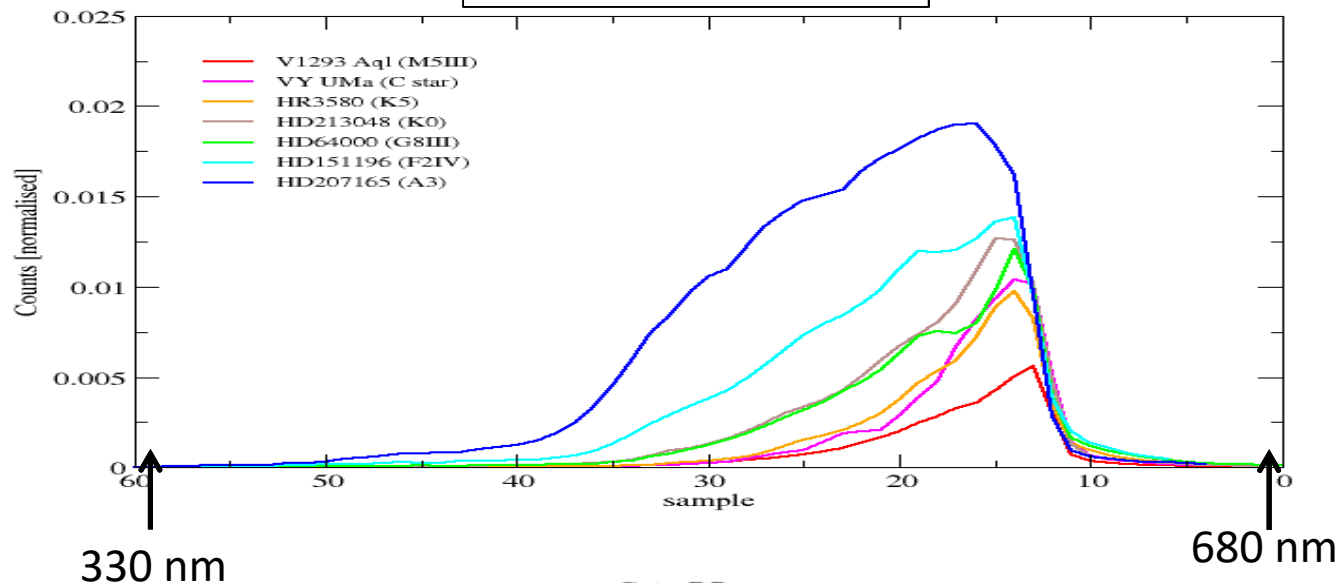
A_V, E_{B-V}

$T_{\text{eff}},$
 $\log g$
 $[\text{Fe}/\text{H}]$
 $[\alpha/\text{Fe}]$

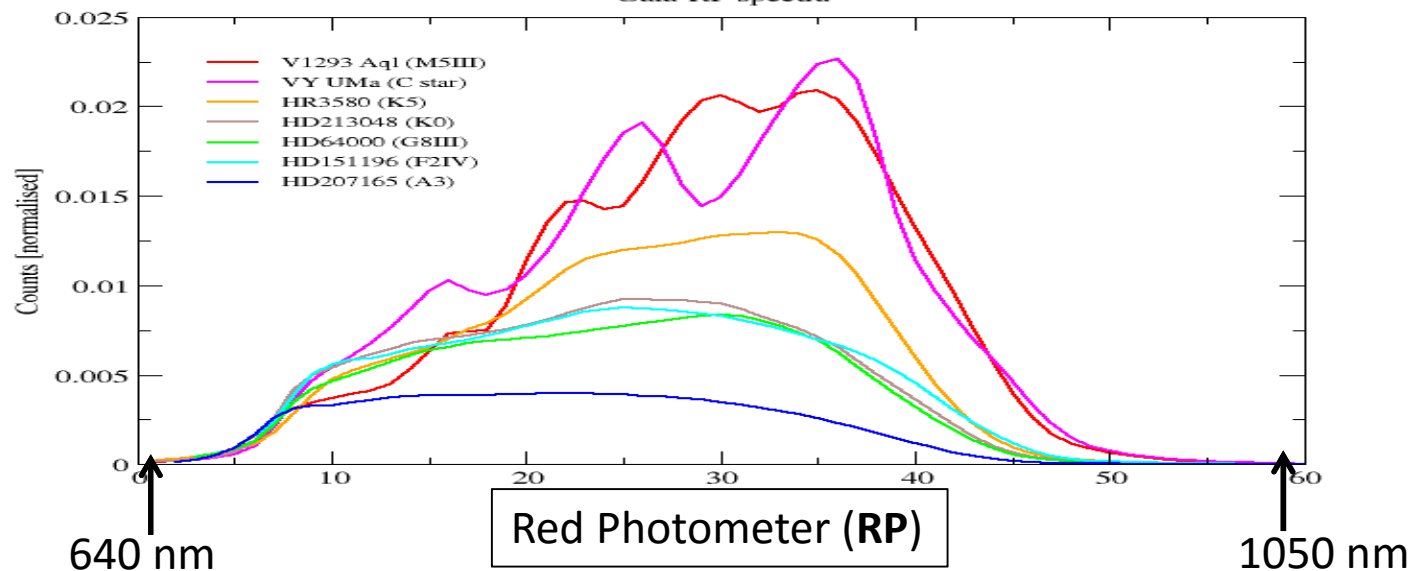
variability

Age

Blue Photometer (BP)



Gaia-RP spectra



Red Photometer (RP)

T_{eff} , $\log g$, $[\text{Fe}/\text{H}]$, $[\alpha/\text{Fe}]$: *RVS*

α , δ , π

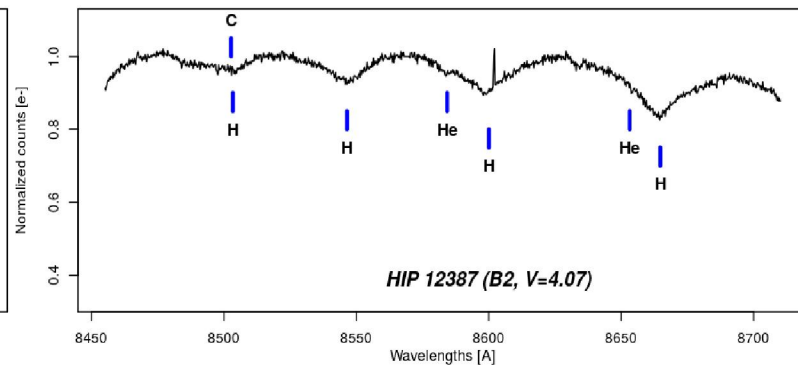
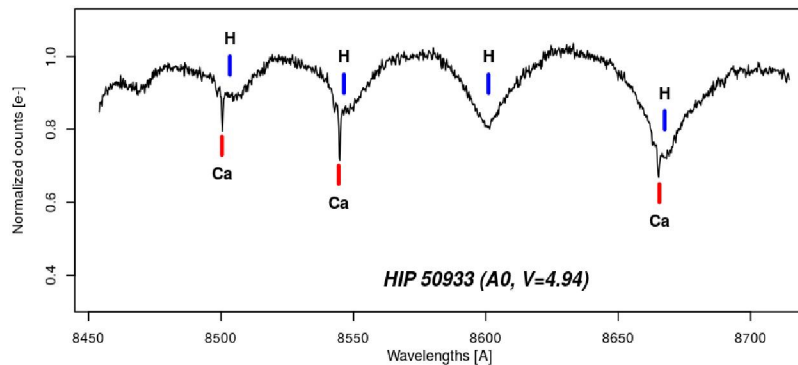
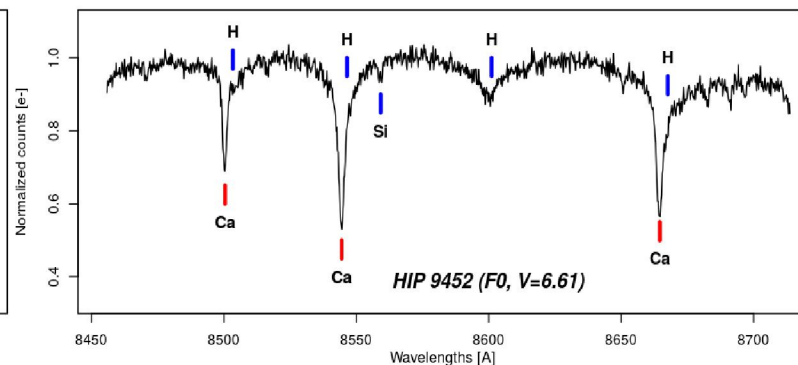
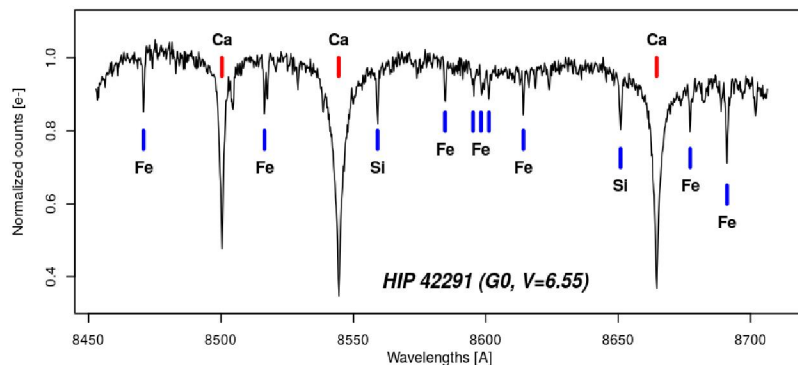
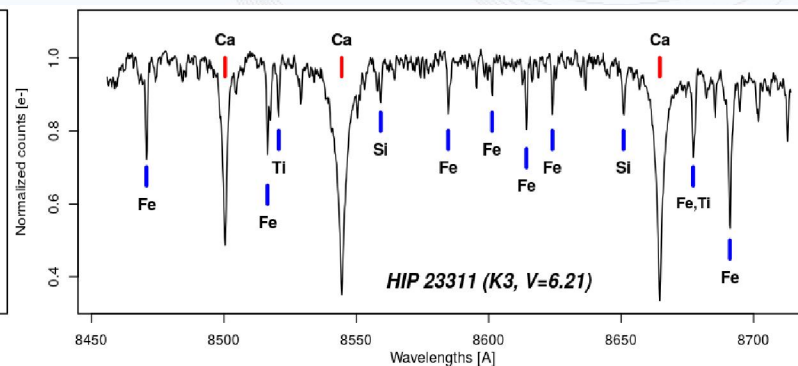
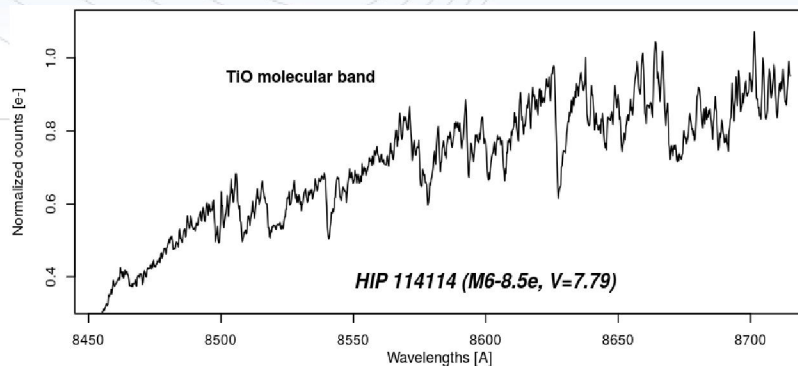
μ_{α} , μ_{δ} , V_R

A_V , E_{B-V}

T_{eff} ,
 $\log g$
 $[\text{Fe}/\text{H}]$
 $[\alpha/\text{Fe}]$

variability

Age



<http://www.cosmos.esa.int/web/gaia>

Credits: ESA/Gaia/DPAC/
Olivier Marchal, David Katz

Teff, logg, [Fe/H], [α/Fe]: precisions

The Gaia astrophysical parameters inference system (Apsis)

Bailer-Jones, Andrae, Arcay et al., 2013, A&A, 559, A74

BP/RP internal precision, **G=15**

$\sigma_{\text{Teff}} = 100 \text{ K}$

$\sigma_{\text{logg}} = 0.25$

$\sigma[\text{Fe}/\text{H}] = 0.2 \text{ dex}$

α, δ, π

$\mu_{\alpha}, \mu_{\delta}, V_R$

A_V, E_{B-V}

**Teff,
logg
[Fe/H]
[X/Fe]**

variability

Age

*Stellar parametrization from Gaia **RVS** spectra*

Recio-Blanco, de Laverny, Allende-Prieto et al., 2016, A&A, 585, A93

Estimated end-of-mission internal precision (extract from Table 5):

	σ_{Teff}		σ_{logg}		$\sigma[\text{Fe}/\text{H}]$		$\sigma[\alpha/\text{Fe}]$	
Grvs	10.3	12.6	10.3	12.6	10.3	12.6	10.3	12.6
K-giant (solar met.)	21	147	0.06	0.29	0.02	0.12	0.02	0.10
K-giant (metal-poor)	65	289	0.13	0.34	0.07	0.25	0.05	0.19

Plato

Teff & Log g seismic

$\sigma_{\text{logg}} = 0.015 \text{ dex}$

(R. Peralta PhD thesis – Corot/Kepler data)

Variability

Credits: ESA/Gaia/DPAC/CU5/DPCI/CU7/INAF-OABO/INAF-OACn Gisella Clementini, Vincenzo Ripepi, Silvio Leccia, Laurent eyer, Lorenzo Rimoldini, Isabelle Lecoeur-Taibi, Nami Mowlavi, Dafydd Eavns, Geneva CU7/DPCG and the whole CU7 team.

α, δ, π

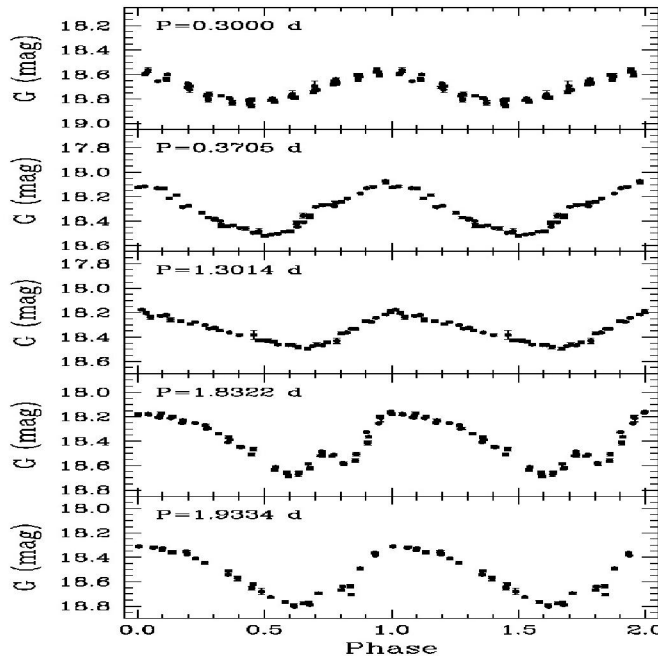
$\mu_\alpha, \mu_\delta, V_R$

A_V, E_{B-V}

$T_{\text{eff}}, \log g$
[Fe/H]
[X/Fe]

variability

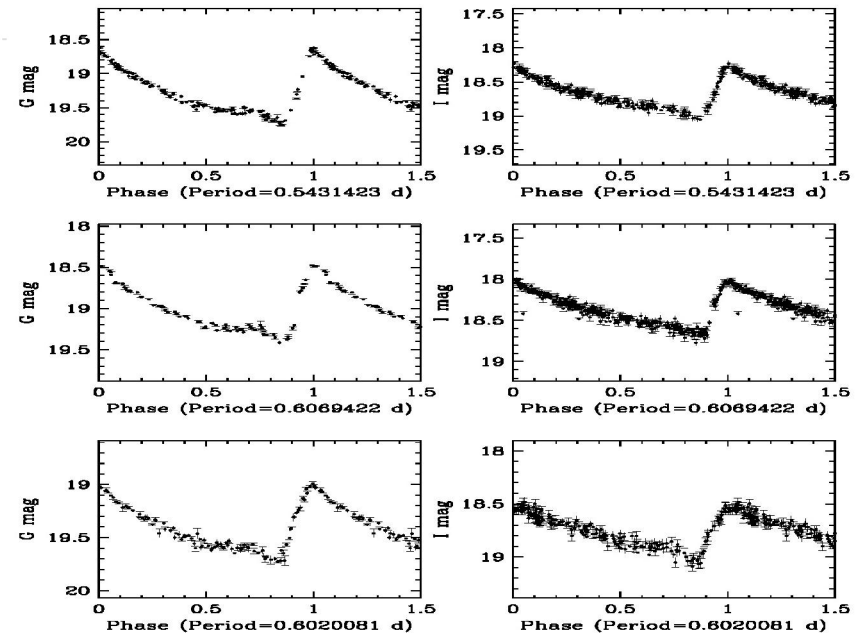
Age



Cepheids in the
Large Magellanic Cloud

Bp/Rp Preliminary processing/calibrations:

- 10-15 mmag per measurement
- 70 epochs (on average in 5 years) + Ecliptic poles



Gaia

Ogle IV

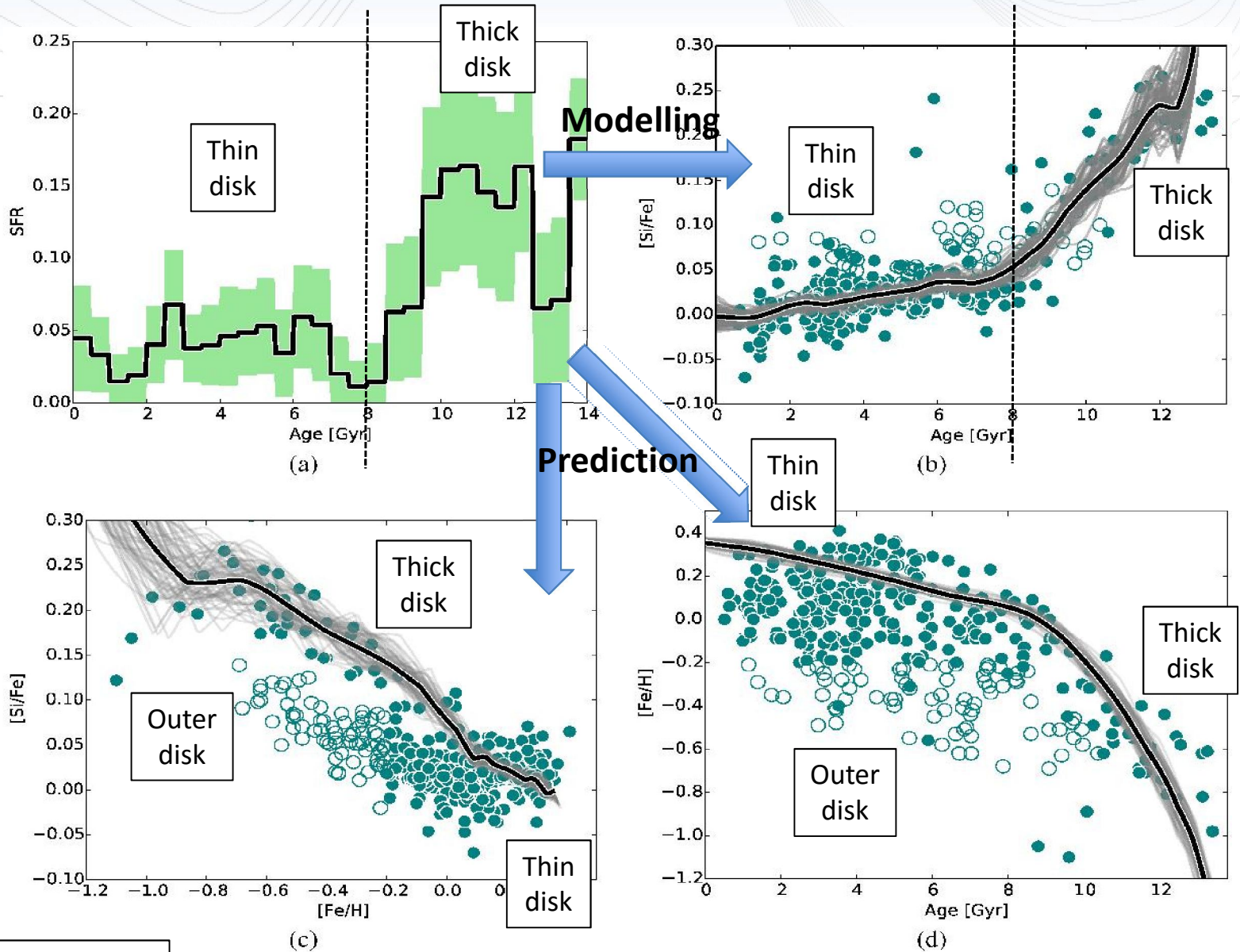
RR-Lyrae in the
Large Magellanic Cloud

Plato

- Much finer time sampling → very detailed oscillation spectra
- Evolutionary stage: separation Red Giant Branch / Red Clump

Age: Milky-way mass assembly & chemical history

α, δ, π
 $\mu_\alpha, \mu_\delta, V_R$
 A_V, E_{B-V}
 $T_{\text{eff}}, \log g$
 $[\text{Fe}/\text{H}]$
 $[\text{X}/\text{Fe}]$
 variability
Age



Ages from isochrones

α, δ, π

$\mu_\alpha, \mu_\delta, V_R$

A_V, E_{B-V}

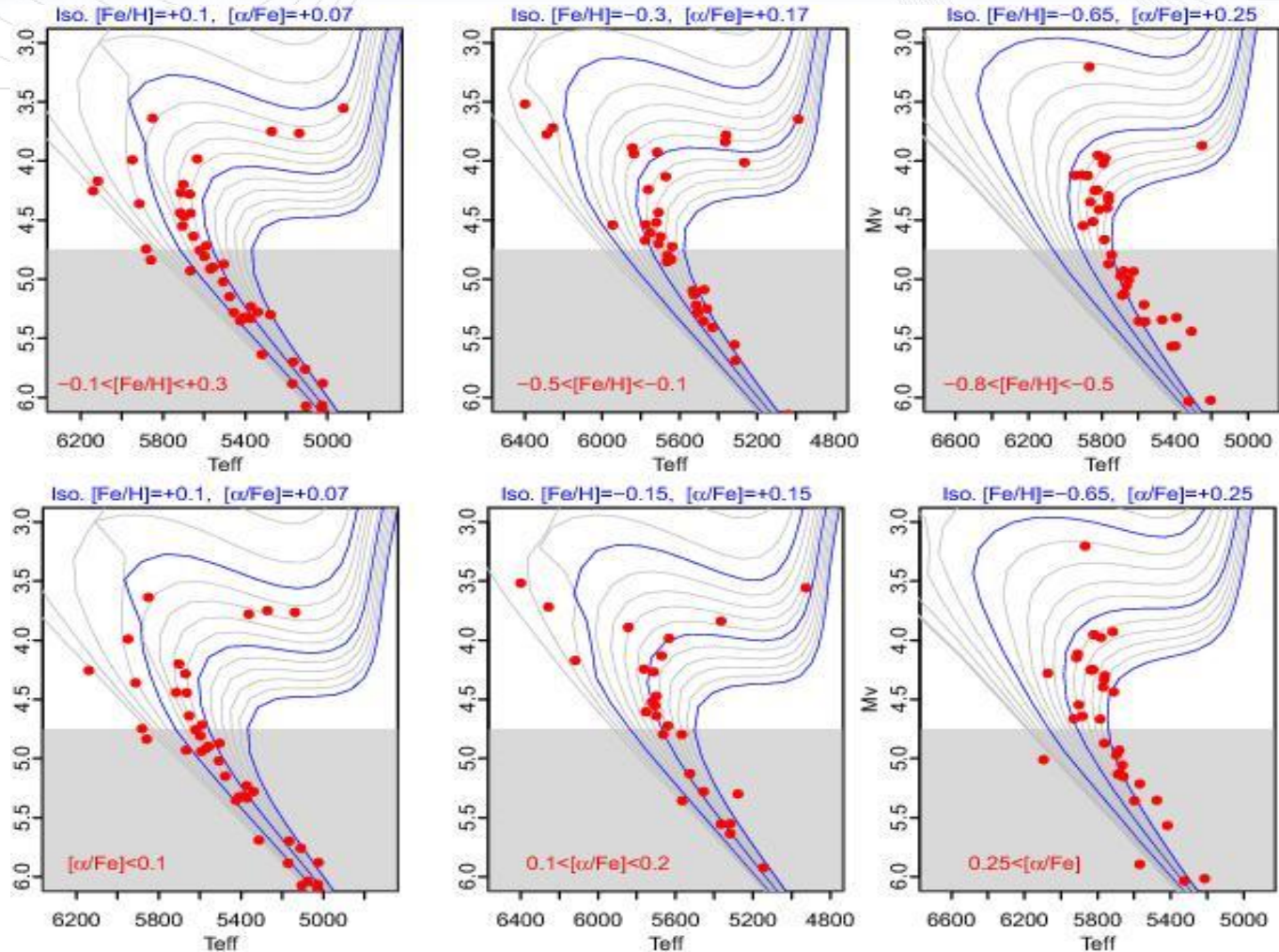
$T_{\text{eff}}, \log g$

$[\text{Fe}/\text{H}]$

$[\text{X}/\text{Fe}]$

variability

Age



- Turn-off/sub-giants: $\sigma_{\text{Age}} \approx 30\%$
- Cool main-sequence / giants: isochrones closely packed

Haywood, Di Matteo, Lehnert et al., 2013, A&A, 560, A109

Ages with Plato

α, δ, π

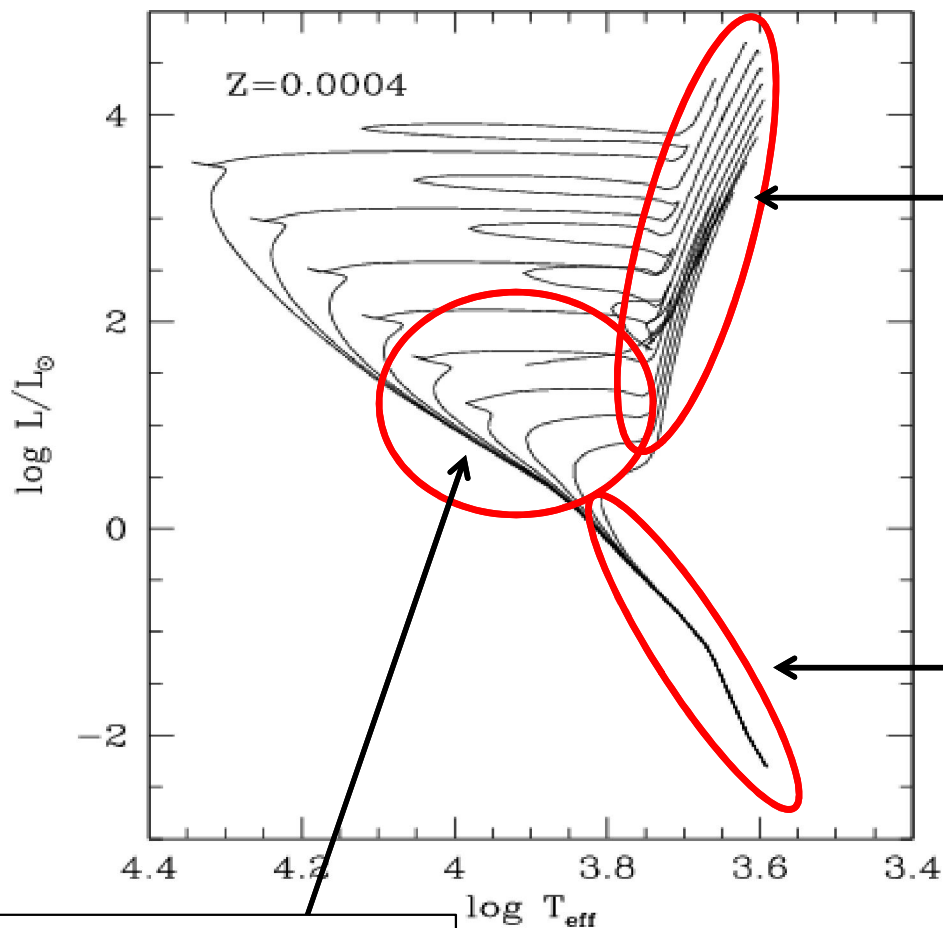
$\mu_\alpha, \mu_\delta, V_R$

A_V, E_{B-V}

$T_{\text{eff}}, \log g$
[Fe/H]
[X/Fe]

variability

Age



Ages

Ages ?

a.)

More precise ages

Radial velocity: e.g. dynamics of accreted satellite

Credits: R. Wyse

α, δ, π

$\mu_\alpha, \mu_\delta,$

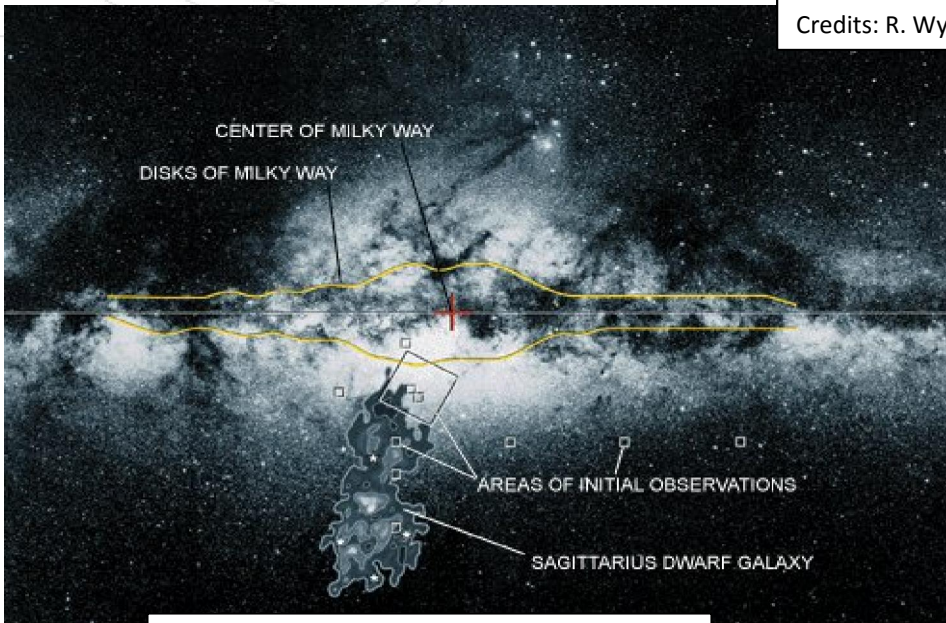
V_R

A_V, E_{B-V}

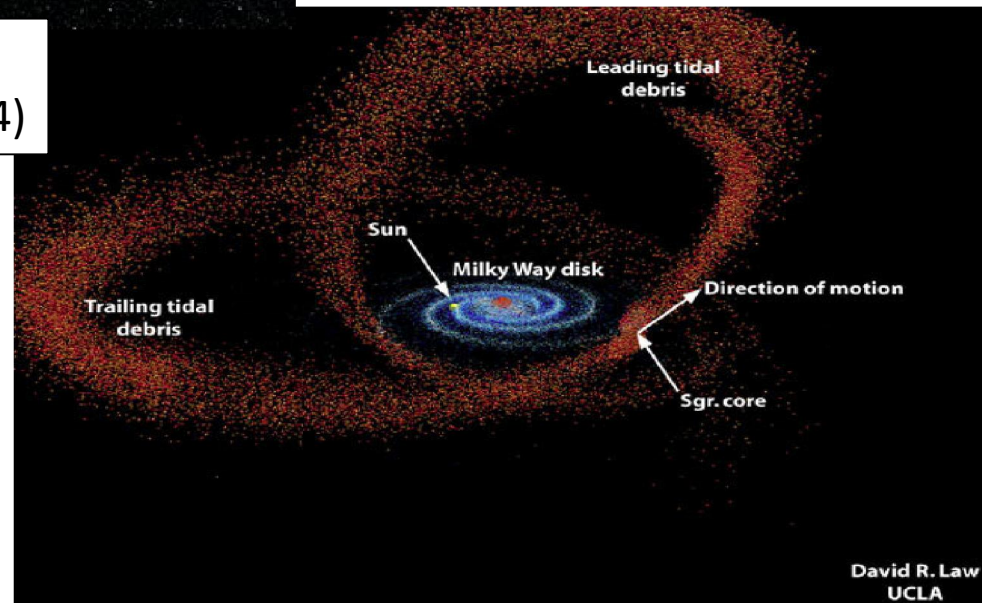
$T_{\text{eff}}, \log g$
 $[\text{Fe}/\text{H}]$
 $[\text{X}/\text{Fe}]$

variability

Age



Sagittarius dwarf galaxy
(discovered: R. Ibata 1994)



A_V, E_{B-V} : e.g., 3D view of the interstellar medium

α, δ, π

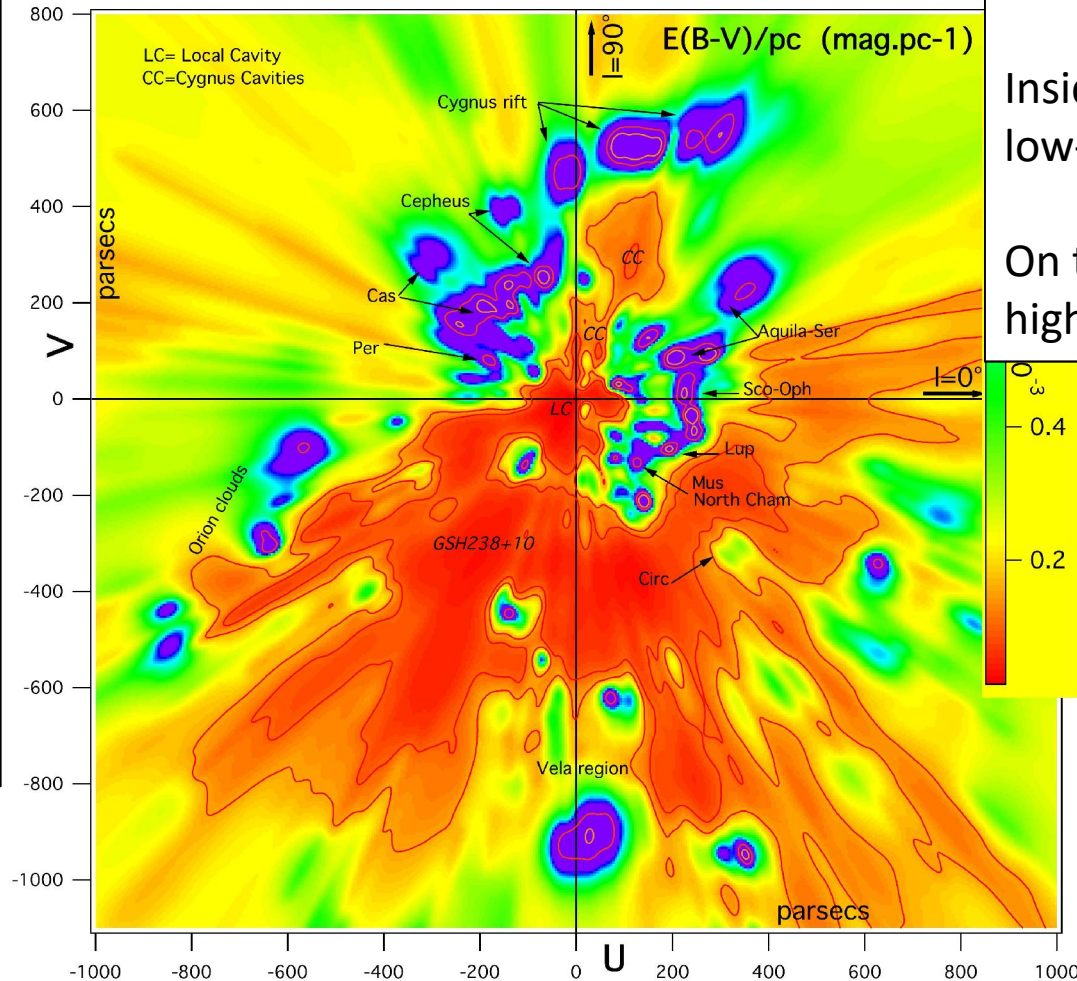
$\mu_\alpha, \mu_\delta, V_R$

$A_V,$
 E_{B-V}

$T_{\text{eff}}, \log g$
 $[\text{Fe}/\text{H}]$
 $[\text{X}/\text{Fe}]$

variability

Age



The interstellar medium absorbs the light of the stars \rightarrow of known distances (Hipparcos)

Inside the local cavity
low-density hot gas ($T \sim 10^6$ K)

On the edges of the local cavity
high density cold gas ($T \sim 10^3$ K)

Lallement, Vergely, Valette et al.,
2013, A&A, 561, A91